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*Herschel: the first science highlights*

LETTER TO THE EDITOR

# The *Herschel* Virgo Cluster Survey

## VI. The far-infrared view of M 87<sup>\*</sup>

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### ABSTRACT

The origin of the far-infrared emission from the nearby radio galaxy M 87 remains a matter of debate. Some studies find evidence of a far-infrared excess due to thermal dust emission, whereas others propose that the far-infrared emission can be explained by synchrotron emission without the need for an additional dust emission component. We present *Herschel* PACS and SPIRE observations of M 87, taken as part of the science demonstration phase observations of the *Herschel* Virgo Cluster Survey. We compare these data with a synchrotron model based on mid-infrared, far-infrared, submm and radio data from the literature to investigate the origin of the far-infrared emission. Both the integrated SED and the *Herschel* surface brightness maps are adequately explained by synchrotron emission. At odds with previous claims, we find no evidence of a diffuse dust component in M 87, which is not unexpected in the harsh X-ray environment of this radio galaxy sitting at the core of the Virgo cluster.

**Key words.** galaxies: individual: M 87 – radiation mechanisms: thermal – radiation mechanisms: non-thermal – infrared: galaxies

### 1. Introduction

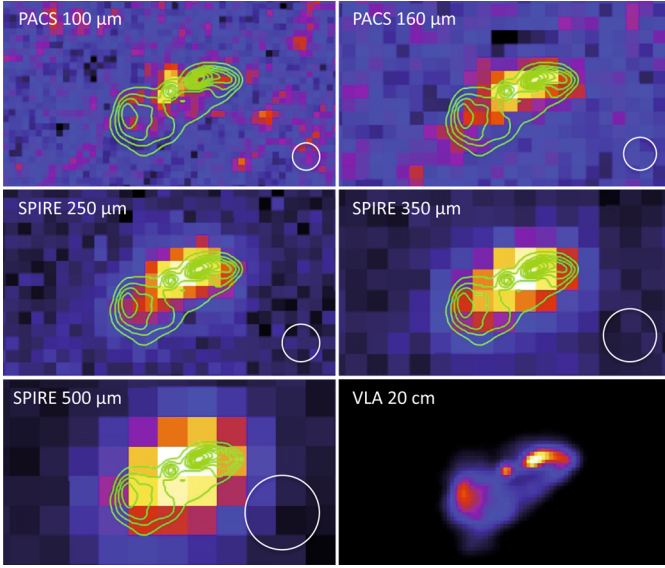
At a distance of 16.7 Mpc (Mei et al. 2007), M 87 is the dominant galaxy of the Virgo cluster. It is one of the nearest radio galaxies and was the first extragalactic X-ray source to be identified. Because of its proximity, many interesting astrophysical phenomena can be studied in more detail in M 87 than in other comparable objects (see e.g. Röser & Meisenheimer 1999, for an overview). Among its many remarkable features is the several billion solar mass supermassive black hole at its centre (Macchetto et al. 1997; Gebhardt & Thomas 2009) and the prominent jet extending from the nucleus, visible throughout the electromagnetic spectrum. The central regions of M 87, in particular the structure of the jet, have been studied and compared intensively at radio, optical, and X-ray wavelengths (e.g. Biretta et al. 1991; Meisenheimer et al. 1996; Böhringer et al. 2001; Perlman et al. 2001; Sparks et al. 2004; Perlman & Wilson 2005; Kovalev et al. 2007; Simionescu et al. 2008; Werner et al. 2010).

Compared to the available information at these wavelengths, our knowledge of M 87 at far-infrared (FIR) wavelengths is rather poor. A controversial issue is the origin of the FIR emission in M 87, i.e., the question of whether the FIR emission is caused entirely by synchrotron emission or whether there is an additional contribution from dust associated with either the global interstellar medium or a nuclear dust component. This question is partly driven by the observation of faint dust features in deep optical images (Sparks et al. 1993; Ferrarese et al. 2006).

Several papers on the FIR emission of M 87 arrive at different conclusions. Perlman et al. (2007) present ground-based Subaru and *Spitzer* IRS spectra of the M 87 nucleus and find evidence of an excess at wavelengths longer than 25  $\mu\text{m}$ , which they attribute to thermal emission from cool dust at a characteristic temperature of some 55 K. This claim is countered by Buson et al. (2009), who present a higher signal-to-noise IRS spectrum of the nucleus. After careful subtraction of a stellar emission template from the mid-infrared spectrum, these authors conclude that the nuclear spectrum can be fully accounted for by optically thin synchrotron emission and that there is little room for dust emission. On a larger scale, Xilouris et al. (2004) present ISOCAM imaging of M 87 and argue that the mid-infrared flux can be attributed to a single synchrotron emission component. Shi et al. (2007) present *Spitzer* IRAC and MIPS imaging of M 87 and find a slight excess in the FIR over a power-law interpolation. They attribute this excess emission to dust emission from the host galaxy. Finally, Tan et al. (2008) observed 1.3 mm continuum emission from the nucleus and jet of M 87 and found that the measured fluxes are generally consistent with synchrotron emission, although they could not rule out a possible nuclear contribution from thermal dust emission.

The recently launched *Herschel* Space Observatory (Pilbratt et al. 2010) offers the possibility to study M 87 at FIR wavelengths in more detail than has been possible to date. The PACS (Poglitsch et al. 2010) and SPIRE (Griffin et al. 2010) instruments combined can produce images over the wavelength range between 70 and 500  $\mu\text{m}$  with unprecedented sensitivity and superior resolution. In this Letter, we present PACS and SPIRE imaging of M 87 at 100, 160, 250, 350, and 500  $\mu\text{m}$ , taken

<sup>\*</sup> *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.



**Fig. 1.** The *Herschel* view of the central regions of M87. The bottom right image is a VLA 20 cm image from the FIRST survey. The 20 cm radio contours have been overlaid on the *Herschel* images. The field of view of all images is  $160'' \times 90''$ , beam sizes are indicated in the bottom right corner.

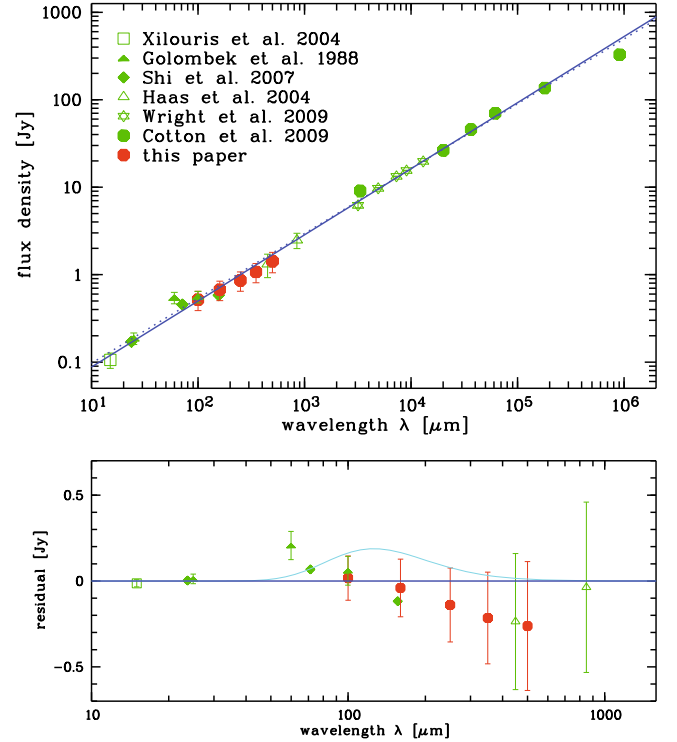
as part of the science demonstration phase (SDP) observations of the *Herschel* Virgo Cluster Survey (HeViCS<sup>1</sup>). We combine these observations with mid-infrared, FIR, submm, and radio data from the literature to investigate the level and the origin of the FIR emission in M87. In Sect. 2, we present the observations and data reduction, Sect. 3 describes the analysis of the data and in Sect. 4 we present our conclusions.

## 2. Observations and data reduction

We observed M87 on 29 November 2009 with PACS and SPIRE as part of the HeViCS SDP observations. The HeViCS SDP field covers a  $4 \times 4 \text{ deg}^2$  field at the centre of the Virgo cluster, roughly centred on M87. It was scanned with a  $60''/\text{s}$  scanning speed in nominal and orthogonal directions. Data were gathered simultaneously in the green and red PACS bands (100 and 160  $\mu\text{m}$ ) and the three SPIRE bands (250, 350, and 500  $\mu\text{m}$ ). The PACS and SPIRE data were reduced using HIPE, with reduction scripts based on the standard reduction pipelines. For more details of the HeViCS SDP data reduction, we refer to Davies et al. (2010). Figure 1 shows the *Herschel* images at the five PACS and SPIRE bands of the central  $160'' \times 90''$  regions of M87, which is clearly detected in all five bands.

## 3. Analysis

The most straightforward, first-order approach to investigating the origin of the FIR emission in M87 is to study its global spectral energy distribution (SED). Table 1 lists the *Herschel* fluxes of M87, together with ISOCAM, IRAS, MIPS, and SCUBA data gathered from the literature (Xilouris et al. 2004; Golombek et al. 1988; Shi et al. 2007; Haas et al. 2004). The top panel in Fig. 2 shows the SED in the infrared-submm-radio region between 15  $\mu\text{m}$  and 100 cm. Apart from the infrared-submm fluxes from Table 1, this plot shows GBT MUSTANG and VLA radio continuum fluxes from Cotton et al. (2009) and the most recent



**Fig. 2.** Top: the global SED of M87 from mid-infrared to radio wavelengths. When no error bars are seen, they are smaller than the symbol size. The solid line in the plot is the best-fit power law of the ISOCAM, IRAS, MIPS, SCUBA, GBT, WMAP, and VLA data; the dotted line has only been fitted to the SCUBA, GBT, WMAP, and VLA data. Bottom: residual between data and the best-fit synchrotron model in the infrared-submm wavelength range. The cyan line is a modified black-body model with  $T = 23 \text{ K}$  and  $M_d = 7 \times 10^4 M_\odot$  (see text).

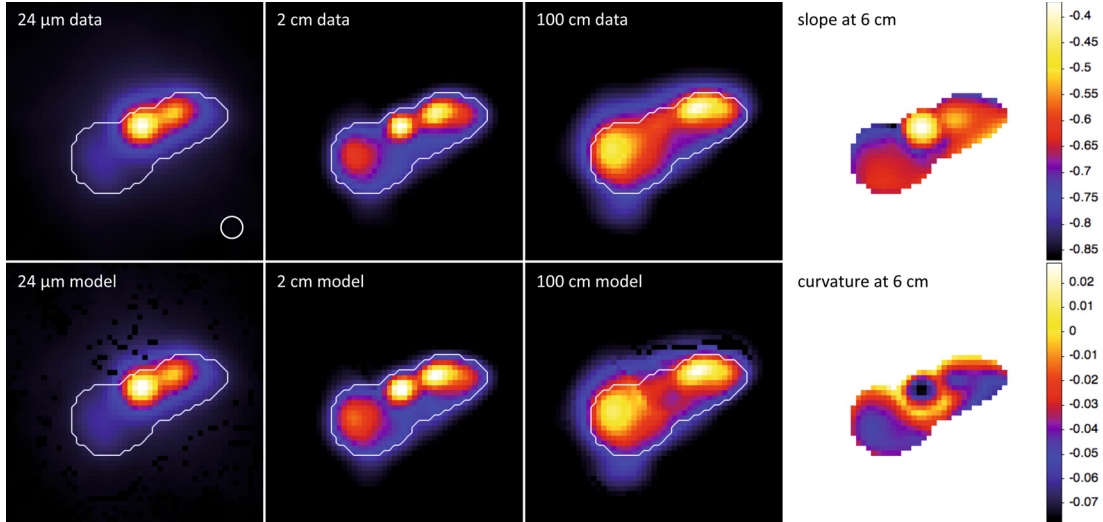
**Table 1.** Integrated fluxes for M87 in the infrared-submm wavelength region between 15 and 1000  $\mu\text{m}$ .

Instrument	$\lambda$ ( $\mu\text{m}$ )	$F_\lambda$ (Jy)	Source
ISOCAM	15	$0.106 \pm 0.021$	Xilouris et al. (2004)
MIPS	24	$0.171 \pm 0.013$	Shi et al. (2007)
IRAS	25	$0.187 \pm 0.028$	Golombek et al. (1988)
IRAS	60	$0.546 \pm 0.082$	Golombek et al. (1988)
MIPS	70	$0.455 \pm 0.009$	Shi et al. (2007)
IRAS	100	$0.559 \pm 0.084$	Golombek et al. (1988)
PACS	100	$0.517 \pm 0.129^a$	this paper
MIPS	160	$0.582 \pm 0.010$	Shi et al. (2007)
PACS	160	$0.673 \pm 0.168^a$	this paper
SPIRE	250	$0.860 \pm 0.215^a$	this paper
SPIRE	350	$1.070 \pm 0.267^a$	this paper
SCUBA	450	$1.320 \pm 0.396$	Haas et al. (2004)
SPIRE	500	$1.430 \pm 0.375^a$	this paper
SCUBA	850	$2.480 \pm 0.496$	Haas et al. (2004)

**Notes.** (a) We assume a conservative uncertainty in the PACS and SPIRE flux densities of 25%, which includes the 15% uncertainty in the absolute flux calibration (Poglitsch et al. 2010; Swinyard et al. 2010), uncertainties introduced by the map-making techniques, and the possible contamination of background objects. Colour corrections were applied, but they turned out to be very small (1.8% for the PACS 100 band and smaller than 1% for the remaining bands).

5-year WMAP fluxes from Wright et al. (2009). The solid line is the best-fit power law for the ISOCAM, IRAS, MIPS, SCUBA, GBT, WMAP, and VLA data and has a slope  $\alpha = -0.76$ ; the dotted line fits only the SCUBA, GBT, WMAP, and VLA data and has a slope  $\alpha = -0.74$ .

<sup>1</sup> More details on HeViCS can be found on <http://www.hevics.org>



**Fig. 3.** A comparison between the data and a synchrotron model. The three images on the *top row* show the observed *Spitzer* image of M 87 at  $24\ \mu\text{m}$  and VLA images at 2 cm and 100 cm, all convolved to the same resolution ( $8.5''$  FWHM,  $2''$  pixels). In each pixel, a synchrotron emission was fit to the *Spitzer*, GBT, and VLA data across the wavelength range between  $24\ \mu\text{m}$  and 100 cm. The solid lines in the images show the region where the data are reliable at all wavelengths: outside this region, not all images could be used. The *bottom row* images show the synchrotron model images corresponding to the images on the top row. The two rightmost images show the slope of the synchrotron fit and the curvature of the fit, both at 6 cm.

The bottom panel in Fig. 2 shows the residual from the best-fit power law in the infrared-submm wavelength region; it is clear that the integrated *Herschel* fluxes are in full agreement with synchrotron radiation. The cyan line in this figure is a modified black-body fitted with  $T = 23\ \text{K}$  and  $M_d = 7 \times 10^4\ M_\odot$ . This temperature is the mean dust equilibrium temperature in the interstellar radiation field of M 87, determined using the SKIRT radiative transfer code (Baes et al. 2003, 2005) and based on the photometry from Kormendy et al. (2009). The dust mass was adjusted to fit the upper limits of the residuals. It is clear that the SED of M 87 is incompatible with dust masses higher than  $10^5\ M_\odot$ .

Although indicative, the analysis of the integrated SED does not definitely identify the origin of the FIR emission in M 87. Approximating the global SED as a single power-law synchrotron model is indeed an oversimplification of the complicated structure of M 87. The bottom-right panel of Fig. 1 shows a high-resolution 20 cm image from the VLA FIRST survey (Becker et al. 1995; White et al. 1997), its contours being superimposed on the *Herschel* images. This 20 cm image identifies three distinct regions of significant synchrotron emission: the nucleus, the jet and associated lobes in the NW region, and the SE lobes. It is well-known that these different components have different spectral indices (e.g., Shi et al. 2007; Cotton et al. 2009) – detailed studies have shown that even within the jet, the spectral index can vary significantly (Biretta et al. 1991; Meisenheimer et al. 1996; Perlman et al. 2001). A spatially resolved analysis of the different regions of M 87 is therefore a more powerful tool for investigating the origin of the FIR emission. Unfortunately, a full spatially resolved analysis of M 87 using all *Herschel* bands is impractical. On the one hand, the PACS PSF is smeared significantly due to the rapid scan speed and the PACS images have a limited signal-to-noise ratio (hereafter S/N). The SPIRE images on the other hand have higher S/N, but the beam size is relatively large for a full spatially resolved analysis of the different components. The M 87 nucleus, jet, and associated NW lobes are discernible in the three SPIRE bands, but the identification of faint emission in the SE lobes is difficult, particularly in the  $500\ \mu\text{m}$  band. We therefore concentrated our efforts on the

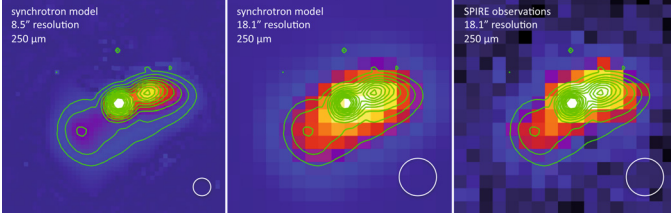
$250\ \mu\text{m}$  image, which provides the optimal compromise between S/N and spatial resolution.

Our analysis consisted of constructing a synchrotron model for the central regions of M 87, based on the available ancillary images with sufficient field-of-view, resolution, and surface brightness sensitivity. We use the new 90 GHz radio continuum maps taken with the MUSTANG bolometer array presented by Cotton et al. (2009). These data have a resolution of  $8''.5$  FWHM, hence closely match our *Herschel* data. The same authors also present archival VLA data at 15, 8.2, 4.9, 1.6, and 0.3 GHz of similar resolution, which were also used in our analysis. An image at 23 GHz was also available in the VLA archive, but it was not used in our analysis because it may be insensitive to extended structure because of a lack of short interferometer baselines. Finally, a *Spitzer* MIPS image at  $24\ \mu\text{m}$  was extracted from the archive and reduced using the MIPS Data Analysis Tools (Gordon et al. 2005), as described in Young et al. (2009). As noted by Shi et al. (2007), a faint extended halo is visible in the  $24\ \mu\text{m}$  image, which is probably caused by stellar emission and/or circumstellar dust. Because of the complexity of the image, we did not attempt to subtract this faint emission from the map; since the flux density in the central regions of M 87 is strongly dominated by non-thermal emission from the nucleus, jet, and lobes, we are confident that this decision does not affect our results.

As a first step, we convolved the available images to the same resolution ( $8''.5$  FWHM), shifted them to the same astrometry ( $2''$  pixel scale), and converted them to similar surface brightness units ( $\text{MJy sr}^{-1}$ ). The result is a data cube with seven points in the wavelength dimension covering an impressive wavelength range of nearly 5 orders of magnitude. Following the same strategy as Cotton et al. (2009), we fitted a second-order polynomial synchrotron model to each pixel of this data cube. The fits were performed with the MPFIT robust non-linear least squares curve fitting library in IDL (Markwardt 2009).

Figure 3 shows a comparison between the data and the resulting synchrotron models at  $24\ \mu\text{m}$ , 2 cm, and 100 cm. The solid white line in these figures borders the spatial region where reliable surface brightnesses were available at all seven





**Fig. 4.** A comparison between the synchrotron model image and the observed image at  $250\ \mu\text{m}$ . The left panel shows the synchrotron image at the model resolution ( $8.5''$  FWHM,  $2''$  pixels), the central panel shows the same model convolved to the SPIRE  $250\ \mu\text{m}$  beam and pixel size ( $18.1''$  FWHM,  $6''$  pixels). The right panel shows the observed SPIRE  $250\ \mu\text{m}$  image. In all panels, the green lines are the contours of the synchrotron model at the model resolution.

wavelengths. Outside this region, spectral fits were still made, but typically used fewer data points, e.g., because one of the radio images contained negative flux at the corresponding wavelengths. By inspecting the images on the top row, it is obvious that the synchrotron emission is not homogeneous in the different regions: the short wavelength radiation is dominated by the central source, whereas the jet and lobes dominate at the longer wavelengths. Comparing the top row panels with the corresponding bottom row panels, however, one can clearly see that, in spite of this different structure, our simple synchrotron model can reproduce the observed images very well over the entire wavelength range. The right-hand panels of Fig. 3 show the slope and the curvature of the synchrotron model at the reference wavelength of 6 cm. The slope varies from  $-0.4$  in the nuclear region to  $-0.8$ , and the curvature is modest ranging from slightly positive in the very faint regions to about  $-0.06$  in nucleus, jet, and lobes. These values are consistent with the values found by Cotton et al. (2009), who limited their fits to the radio regime between 3.3 mm and 100 cm.

For the next step in our analysis, we used our synchrotron model to predict the emission of M 87 at  $250\ \mu\text{m}$ . The left panel of Fig. 4 shows the synchrotron model prediction at  $250\ \mu\text{m}$  at the model resolution. By far the brightest peak of the model is located at the position of the nucleus; the jet is clearly visible as a second bright component, whereas the SE lobe is visible as an extended, low surface brightness region with a peak surface brightness that is almost five times fainter than the nucleus ( $23\ \text{MJy sr}^{-1}$  versus  $110\ \text{MJy sr}^{-1}$ ). When we convolve this synchrotron model image with the SPIRE  $250\ \mu\text{m}$  beam and rebin it to match the observed SPIRE image astrometry, the three different components corresponding to nucleus, jet, and SE lobes merge into a single extended structure with one elongated peak  $4''$  to the west of the nucleus. Comparing the central and right panels of Fig. 4, we see that the synchrotron model is capable of explaining the observed SPIRE  $250\ \mu\text{m}$  image satisfactorily.

#### 4. Conclusion

For both the integrated SED and the SPIRE  $250\ \mu\text{m}$  map, we have found that synchrotron emission is an adequate explanation of the FIR emission. We do not detect a FIR excess that cannot be explained by the synchrotron model. In particular, we have no reason to invoke the presence of smooth dust emission associated with the galaxy interstellar medium, as advocated by Shi et al. (2007). For a dust temperature of 23 K, which is the expected equilibrium temperature in the interstellar radiation field of M 87, we find an upper limit to the dust mass of  $7 \times 10^4\ M_{\odot}$ .

Our result agrees with the analysis of the nuclear emission by Buson et al. (2009). Clemens et al. (2010) discuss the lifetimes of interstellar dust grains in elliptical galaxies in the Virgo cluster based on *Herschel* observations and find an upper limit to the amorphous silicate grain survival time of less than 46 million years. Given that M 87 is a luminous X-ray source, the absence of a substantial dust component is not a surprise. A low dust content is also in agreement with the non-detection of cool molecular gas (Salomé & Combes 2008; Tan et al. 2008) and the non-detection of significant intrinsic absorption in the X-ray spectra of M 87 (Böhringer et al. 2002). Our conclusion is that, seen from the FIR point of view, M 87 is a passive object with a central radio source emitting synchrotron emission, without a substantial diffuse dust component.

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